

Aspects of X-ray data analysis for accreting compact objects: theory and results

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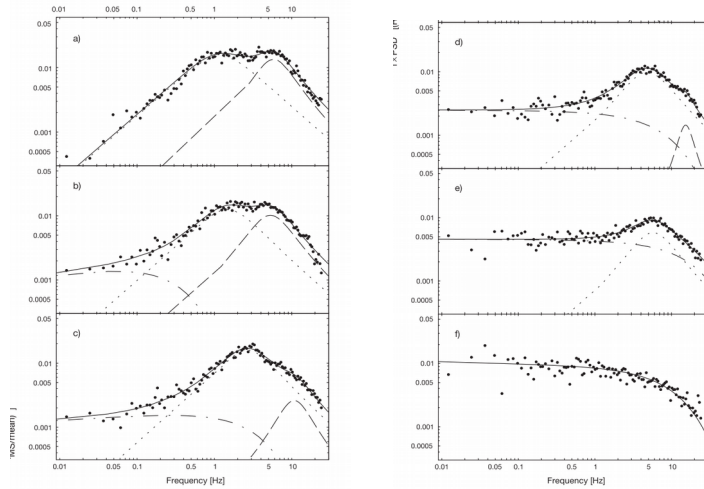
Nicolaus Copernicus Astronomical Center, Polish Academy of Sciences

Lecture 6, 25.11.2025

PhD lecture series, 2025, fall semester

Previous lecture

X-ray variability



Correlations between parameters

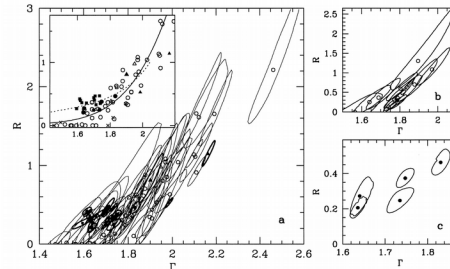
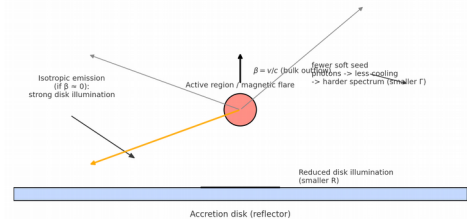
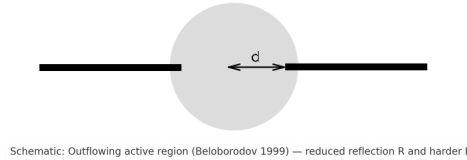
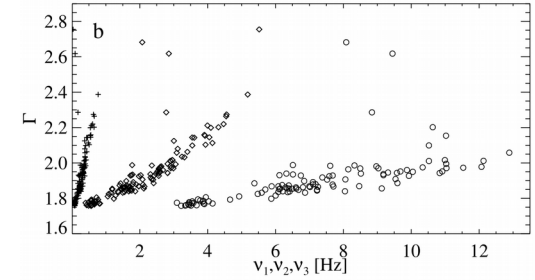


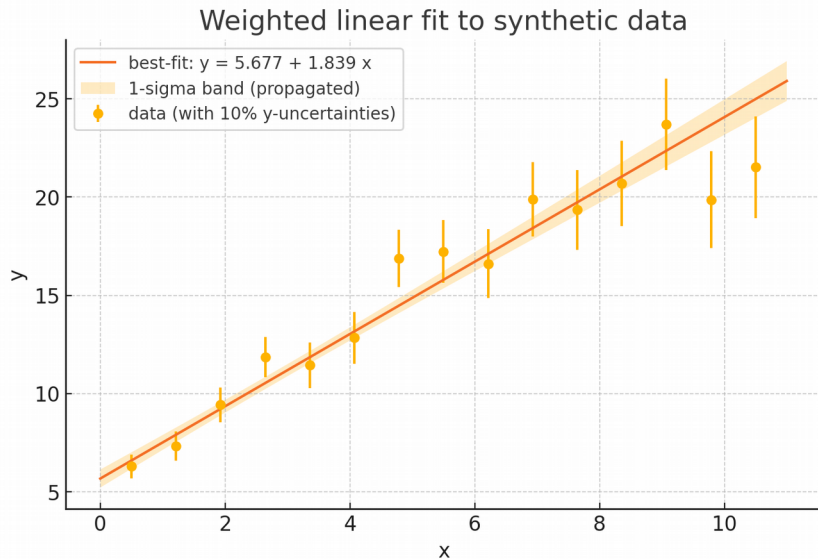
Figure 1. The R/Γ correlation in Seyfert and X-ray binaries in the hard state. Panel (a) shows the data and models (curves in the inset; see Sections 2 and 4, respectively). Examples of the correlation for NGC 5548 and GX 339-4 are shown in panels (b) and (c) respectively.

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Evolution of PDS: from a multiple-Lorentzian to a cutoff power law

Data modelling



$\chi^2 = 11.5$ for 13 d.o.f (15 data points – 2 parameters)

Chat GPT

$$\chi^2(\theta) = \sum_i \frac{[y_i - f(x_i, \theta)]^2}{\sigma_i^2}$$

θ – parameters of the model

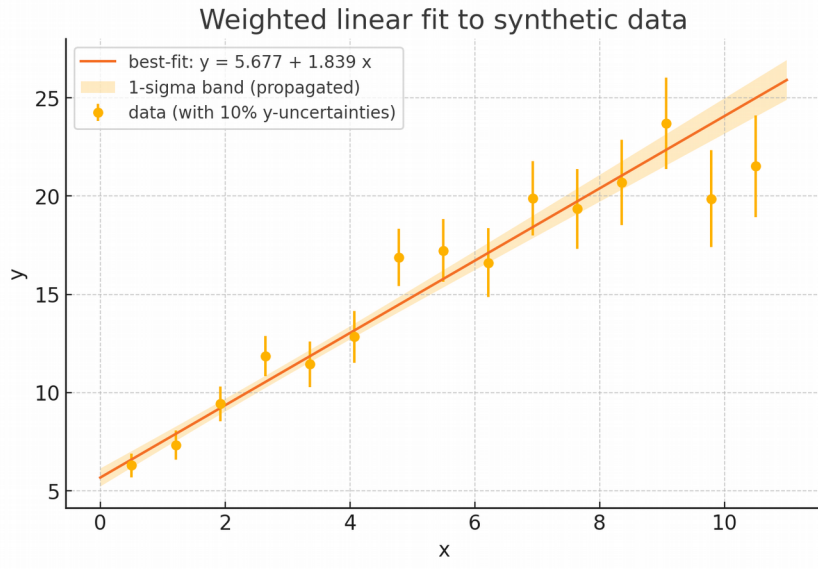
We want to find the parameters of the model function that minimize χ^2 .

Simple case: $f(x) = ax + b$

In this simple case there exist analytical formulae for a and b .

In more complex cases (non-linear model functions) we need to use numerical algorithms to find the minimum of χ^2 .

Data modelling



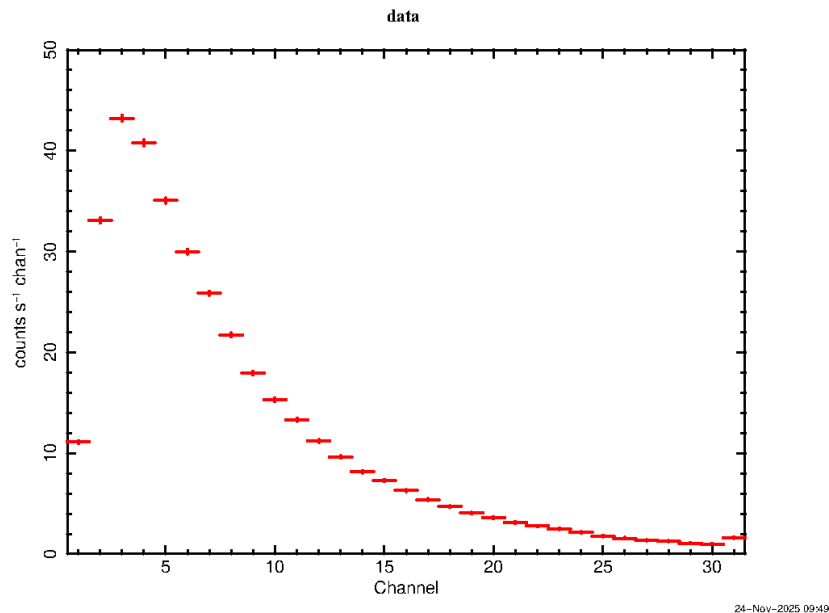
The error bars of the data points represent uncertainty of the location of the real y-value for a given x.

Therefore, the parameters of our model, a and b , are also uncertain. The fitting procedure should also give us these uncertainties.

In the simple case the uncertainties are also given by analytical formulae. **For more complex cases, the procedures to estimate the uncertainties are more complex.**

Meaning of the $1\text{-}\sigma$ band: imagine making N measurements, so we have N sets of points. The model is fitted to the data and $1\text{-}\sigma$ band is computed. For a given x , the $1\text{-}\sigma$ band contains the true value 68% of the times.

X-ray data

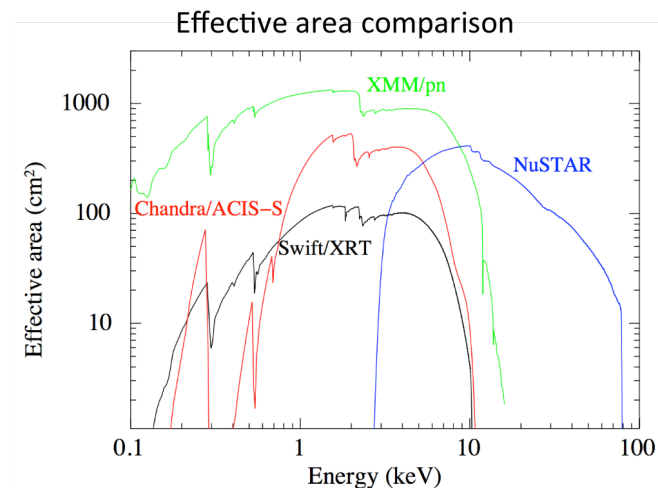


Ginga (1990-ies) – effective area of 4000 cm²,
poor energy resolution

Response function (matrix) – redistribution of
incoming photons in energy

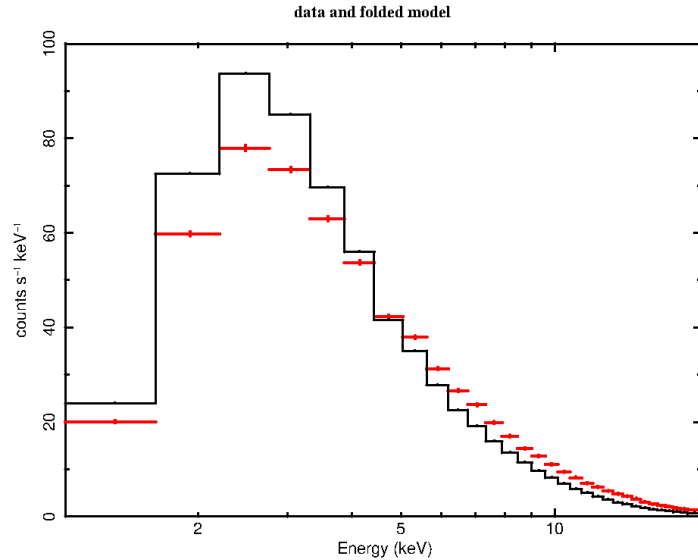
Effective area – energy dependent

Energy channels

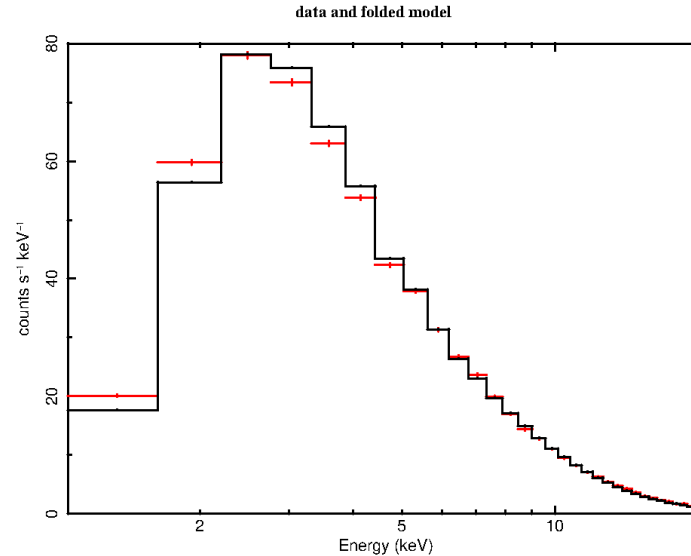


Modelling X-ray data

GS 1124-68 – Nova Muscae – a soft X-ray transient



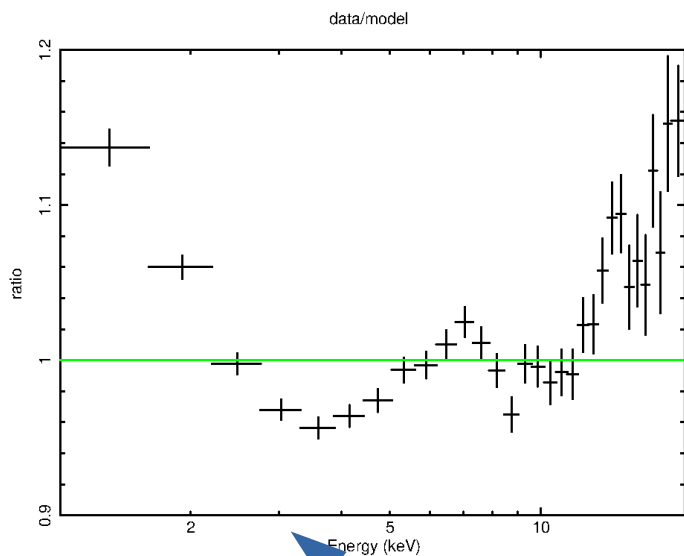
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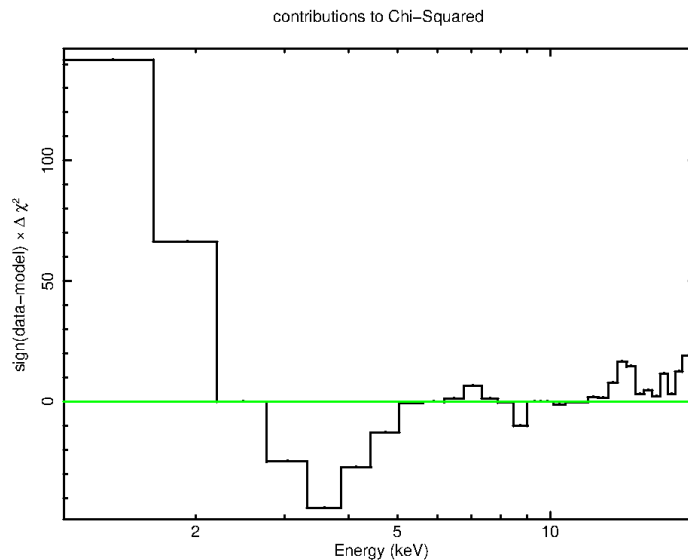
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Initial guess of the model and parameters → not optimum description of data → fitting (minimizing χ^2) → best fit

Modelling X-ray data



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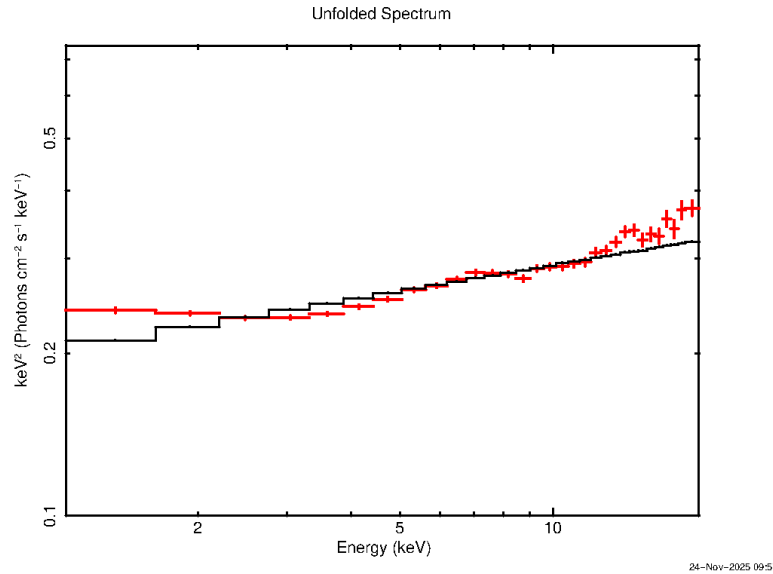


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Diagnostics of the fit results:
data-to-model ratios,
contributions to χ^2

This fit is not so good, significant residuals are seen, χ^2 large.

Modelling X-ray data



Unfolded spectrum: how the real data would look like, assuming the best-fit model.

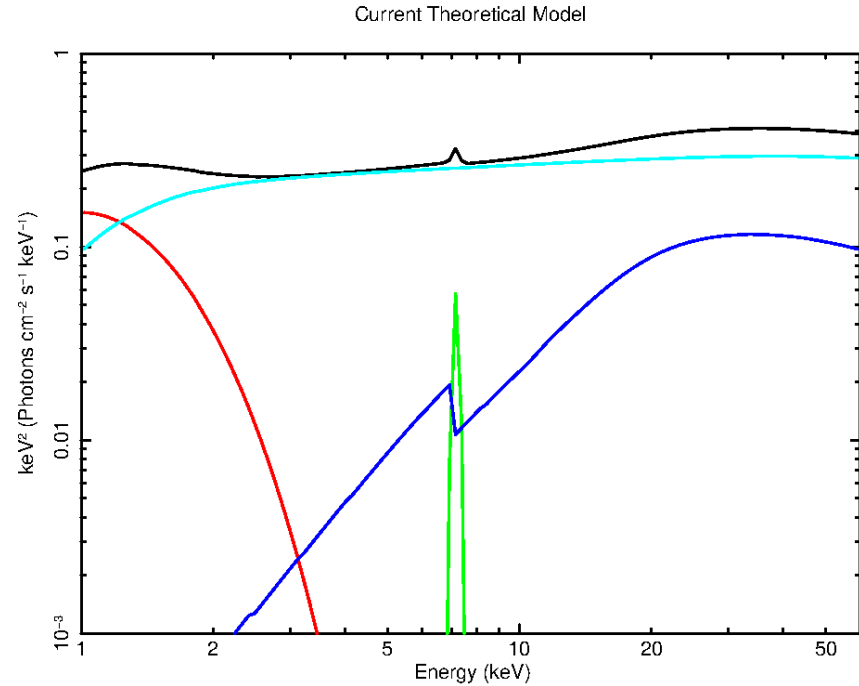
Key aspect: it is not possible to uniquely determine the real data (before they reach the detector), the unfolded spectrum depends on the model.

Finding a good fit

Need a better model. Using previous knowledge, physics, intuition, fit residuals, trial and error...

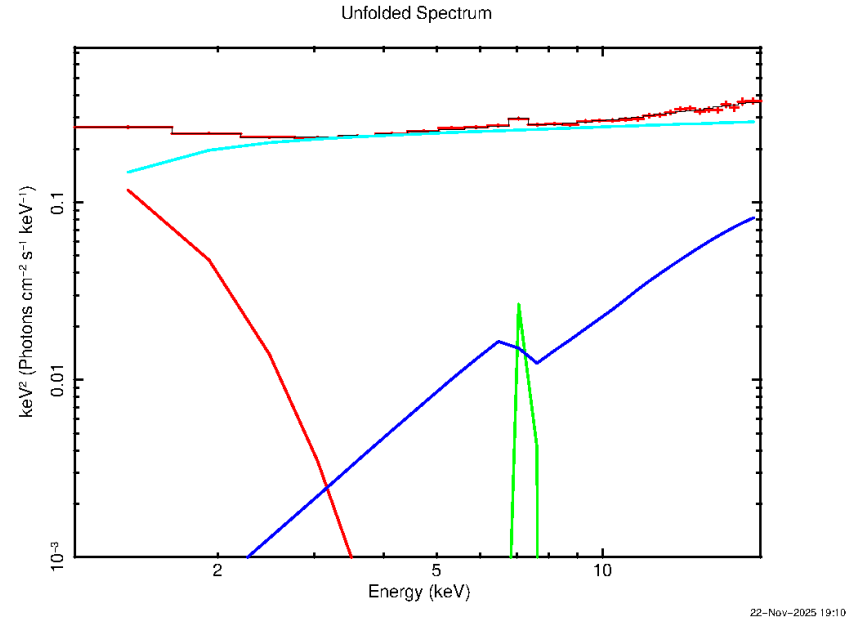
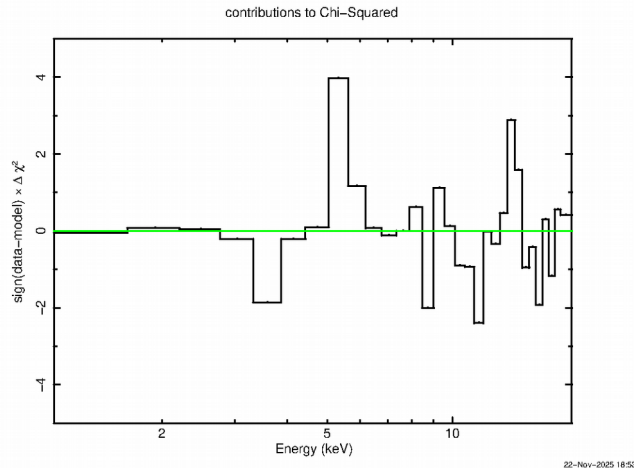
This model components are: disk blackbody emission, thermal Comptonization component, reflected/reprocessed component, with the Fe K_α line.

Best fit of this model is good: $\chi^2 = 27/24$ d.o.f.



Finding a good fit

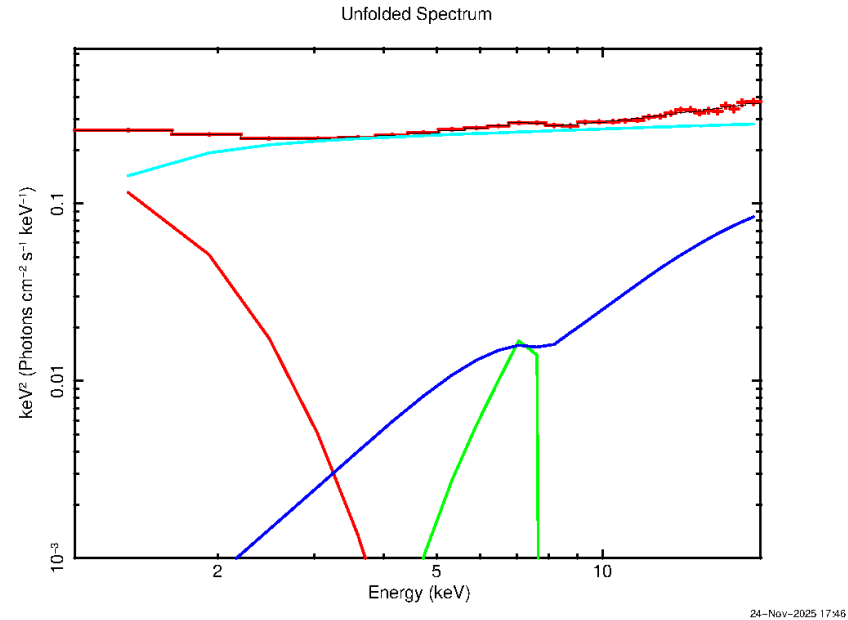
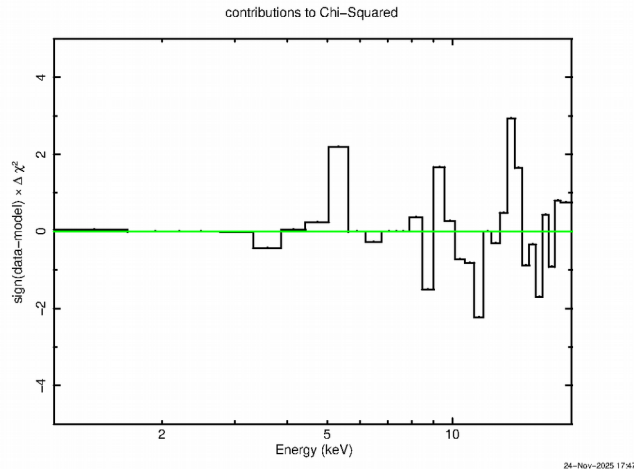
Best fit of this model is good: $\chi^2 = 27/24$ d.o.f.



This is a good fit, although some residuals remain. They suggest a way to improve the fit even more, by applying relativistic effects.

Finding a better fit

Best fit of this model is good: $\chi^2 = 22/23$ d.o.f.



A better fit, suggesting that the emission is coming from a rotating disk

Properties of the good fit

```
=====
Model wabs<1>(diskbb<2> + gaussian<3> + pexrav<4> + nthComp<5>) Source No.: 1 Active/On
Model Model Component Parameter Unit Value
par comp
  1 1 wabs nH 10^22 0.160000 frozen
  2 2 diskbb Tin keV 0.279110 +/- 1.67868E-02
  3 2 diskbb norm 6610.92 +/- 2501.30
  4 3 gaussian LineE keV 7.16468 +/- 0.131805
  5 3 gaussian Sigma keV 0.100000 frozen
  6 3 gaussian norm 3.44840E-04 +/- 6.45261E-05
  7 4 pexrav PhoIndex 1.89245 = p15
  8 4 pexrav foldE keV 300.000 frozen
  9 4 pexrav rel_refl -0.925049 +/- 0.150065
10 4 pexrav Redshift 0.0 frozen
11 4 pexrav abund 1.00000 frozen
12 4 pexrav Fe_abund 1.00000 frozen
13 4 pexrav cosIncl 0.500000 frozen
14 4 pexrav norm 0.139455 = p20
15 5 nthComp Gamma 1.89245 +/- 1.57632E-02
16 5 nthComp kT_e keV 100.000 frozen
17 5 nthComp kT_bb keV 0.279110 = p2
18 5 nthComp inp_type 0/1 0.0 frozen
19 5 nthComp Redshift 0.0 frozen
20 5 nthComp norm 0.139455 +/- 9.89552E-03

XSPEC12>sho fit

Fit statistic : Chi-Squared 26.99 using 31 bins.

Test statistic : Chi-Squared 26.99 using 31 bins.
Null hypothesis probability of 3.05e-01 with 24 degrees of freedom
```

Properties of the good fit

How to estimate the uncertainties of the fitted parameters?

Use χ^2 as a function of the parameters.

Confidence limits can be estimated from $\Delta\chi^2$ relative to the minimum. For one parameter $\Delta\chi^2=2.71$ corresponds to 90% confidence limit

```
XSPEC12>steppar 15 1.83 1.93 10
```

Chi-Squared	Delta Chi-Squared		Gamma 15
43.846	16.854	0	1.83
38.77	11.778	1	1.84
34.613	7.6208	2	1.85
31.38	4.3879	3	1.86
29.056	2.0637	4	1.87
27.618	0.62591	5	1.88
27.018	0.025536	6	1.89
27.21	0.21769	7	1.9
28.175	1.1829	8	1.91
29.886	2.8937	9	1.92
32.326	5.3341	10	1.93

Properties of the good fit

The amplitude of reflected component

```
XSPEC12>steppar 9 -1.5 -0.5 10
```

Chi-Squared	Delta Chi-Squared		rel_refl 9
38.644	11.652	0	-1.5
35.219	8.2273	1	-1.4
32.288	5.296	2	-1.3
29.933	2.9407	3	-1.2
28.218	1.2258	4	-1.1
27.222	0.23023	5	-1
27.021	0.028399	6	-0.9
27.685	0.6931	7	-0.8
29.301	2.309	8	-0.7
31.952	4.9598	9	-0.6
35.71	8.7179	10	-0.5

Contours of χ^2 as function of spectral index and amplitude of reflection

